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ON THE OBSERVED DISTRIBUTIONS OF BLACK HOLE MASSES AND EDDINGTON RATIOS FROM RADIATION PRESSURE CORRECTED VIRIAL INDICATORS

ALESSANDRO MARCONI¹, DAVID J. AXON², ROBERTO MAIOLINO³, TOHRU NAGAO⁴, PAOLA PIETRINI¹, GUIDO RISALITI^{5,6},
ANDREW ROBINSON², AND GUIDETTA TORRICELLI⁵

¹ Dipartimento di Astronomia e Scienza dello Spazio, Università degli Studi di Firenze, Largo E. Fermi 2, 50125 Firenze, Italy; marconi@arcetri.astro.it

² Physics Department, Rochester Institute of Technology, 85 Lomb Memorial Drive, Rochester, NY 14623, USA

³ INAF-Osservatorio Astronomico di Roma, Via Frascati 33, I-00040 Monte Porzio Catone, Italy

⁴ Research Center for Space and Cosmic Evolution, Ehime University, 2-5 Bunkyo-cho, Matsuyama 790-8577, Japan

⁵ INAF-Osservatorio Astrofisico di Arcetri, Largo E. Fermi 5, 50125 Firenze, Italy

⁶ Harvard-Smithsonian Center for Astrophysics, 60 Garden Street, Cambridge, MA 02138, USA

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ABSTRACT

The application of the virial theorem to the broad-line region (BLR) of active galactic nuclei (AGNs) allows black hole (BH) mass estimates for large samples of objects at all redshifts. In a recent paper, we showed that ionizing radiation pressure onto BLR clouds affects virial BH mass estimates and we provided empirically calibrated corrections. More recently, a new test of the importance of radiation forces has been proposed: the $M_{\text{BH}}-\sigma$ relation has been used to estimate M_{BH} for a sample of type-2 AGNs and virial relations (with and without radiation pressure) for a sample of type-1 AGNs extracted from the same parent population. The observed L/L_{Edd} distribution based on virial BH masses is in good agreement with that based on $M_{\text{BH}}-\sigma$ only if radiation pressure effects are negligible, otherwise significant discrepancies are observed. In this Letter, we investigate the effects of intrinsic dispersions associated with the virial relations providing M_{BH} , and we show that they explain the discrepancies between the observed L/L_{Edd} distributions of type-1 and type-2 AGNs. These discrepancies in the L/L_{Edd} distributions are present regardless of the general importance of radiation forces, which must be negligible only for a small fraction of sources with large L/L_{Edd} . Average radiation pressure corrections should then be applied in virial M_{BH} estimators until their dependence on observed source physical properties has been fully calibrated. Finally, the comparison between M_{BH} and L/L_{Edd} distributions derived from σ -based and virial estimators can constrain the variance of BLR physical properties in AGNs.

Key words: galaxies: active – galaxies: fundamental parameters – galaxies: nuclei – galaxies: Seyfert – quasars: emission lines – radiation mechanisms: general

1. INTRODUCTION

In the last few years, it has become increasingly clear that supermassive black holes (BHs) are an essential element in the evolution of galaxies. The key observational evidence of a link between a BH and its host galaxy is provided by the tight correlations between BH mass and luminosity, mass, velocity dispersion, and surface brightness profile of the host spheroids (e.g., Gebhardt et al. 2000; Ferrarese & Merritt 2000; Marconi & Hunt 2003; Graham & Driver 2007). The link between BH and host galaxy is probably established by the feedback of the accreting BH on the host galaxy itself (e.g., Granato et al. 2004; Di Matteo et al. 2005; Croton et al. 2006, and references therein).

In order to fully understand the implications of BH growth on host galaxy evolution, it is essential to measure BH masses in large samples of galaxies from zero to high redshifts. Since direct BH mass measurements based on stellar and gas kinematics are possible only in the local universe (e.g., Ferrarese & Ford 2005), less direct estimators have been calibrated following a “BH mass ladder” (Peterson 2004). The final rung of this ladder is provided by the virial estimators which allow us to estimate BH masses from the spectra of active galactic nuclei (AGNs) with broad emission lines (type-1): under the assumption that the broad line region (BLR) is gravitationally bound and its motions virialized. The BH mass can be expressed as $M_{\text{BH}} = \tilde{f}V^2 R_{\text{BLR}}/G$, where R_{BLR} is the BLR average distance from the BH, V is the width of the broad emission line, and \tilde{f} is a scaling factor which depends on (unknown) BLR

properties. R_{BLR} can be estimated with the $R_{\text{BLR}}-L$ relation ($R_{\text{BLR}} \propto L^\gamma$; Kaspi et al. 2000; Bentz et al. 2008) leading to $M_{\text{BH}} = fV^2L^\gamma$, where f is calibrated empirically starting from the $M_{\text{BH}}-\sigma$ relation (e.g., Onken et al. 2004; Vestergaard & Peterson 2006).

One of the basic assumptions of reverberation mapping is that the BLR is photoionized (Blandford & McKee 1982) implying that BLR clouds are subject to radiation forces arising from ionizing photon momentum deposition. In a recent paper (Marconi et al. 2008, hereafter M08), we showed that these radiation forces constitute an important physical effect which must be taken into account when computing virial BH masses. We empirically calibrated a radiation pressure correction of the form $M_{\text{BH}} = fV^2L^\gamma + gL$, and we showed that it is consistent with a simple physical model in which BLR clouds are optically thick to ionizing radiation and have average column densities of $N_{\text{H}} \simeq 10^{23} \text{ cm}^{-2}$ toward the ionizing source. This value is remarkably similar to that adopted in standard photoionization models to explain observed BLR spectra.

Recently, Netzer (2009, hereafter N09) proposed a test of the importance of radiation forces on virial BH mass estimates. He selected two large samples of type-2 and type-1 radio-quiet AGNs drawn from the Sloan Digital Sky Survey, and covering the same range of redshift ($0.1 \leq z \leq 0.2$) and continuum luminosity ($10^{42.8} \leq \lambda L_\lambda(5100 \text{ \AA}) \leq 10^{44.8} \text{ erg s}^{-1}$). After eliminating type-2 galaxies classified as LINERS resulted in a final sample composed of 4197 and 1331 in the type-2 and type-1 objects, respectively. By comparing the distributions of

[O III] line luminosities, he concluded that the two samples were extracted from the same parent population. After the sample selection, N09 estimated BH masses for the type-2 sample using the $M_{\text{BH}}-\sigma$ relation in the calibration by Tremaine et al. (2002). For the type-1 sample, he estimated BH masses using both the classical virial relation (e.g., Vestergaard & Peterson 2006) and the one by M09 with the correction for radiation pressure. He then compared the distribution of BH masses and L/L_{Edd} ratios finding significantly different distributions between type-2 and type-1 AGNs under the assumption that radiation pressure significantly affects BLR motions. In particular, while the L/L_{Edd} distribution of type-2 AGNs is broad, bell-shaped, and extends up to $L/L_{\text{Edd}} \sim 1$, the L/L_{Edd} distribution of type-1 AGNs with an important radiation pressure correction is strongly peaked at $L/L_{\text{Edd}} \simeq 0.15$ with a sharp cutoff at $L/L_{\text{Edd}} \sim 0.15$ (see Figure 3 of N09). Conversely, the distribution of type-1 AGNs without the radiation pressure correction is in good agreement with that of type-2 AGNs. N09 then concluded that radiation forces are not affecting BLR motions, which is possible only if BLR clouds have extremely large column densities ($N_{\text{H}} \sim 10^{24} \text{ cm}^{-2}$).

In this Letter, we generalize the conclusions by N09 and we investigate the origin of the discrepancies among the L/L_{Edd} distributions. The observed discrepancies do not imply that radiation forces are not important. They are expected from the physically justified mathematical expression of radiation pressure corrected virial masses, if the intrinsic dispersion associated with such scaling relations is not taken into account. In particular, to obtain the “true” M_{BH} and L/L_{Edd} distributions, one should take into account the dispersion of the scaling parameters used in the virial estimators. We will not address the issue as to whether outflows affect BLR motions (e.g., Chiang & Murray 1996; Kurosawa & Proga 2009) but we will assume that BLR clouds are gravitationally bound.

In Section 2, we discuss why truncated L/L_{Edd} distributions are expected when using virial mass indicators with the correction for radiation pressure. In Section 3, we use Monte Carlo simulations to reproduce the results of N09 and we show that similar observed distributions are expected regardless of the importance of radiation pressure. Finally, we summarize our results and draw our conclusions in Section 4.

2. OBSERVED AND TRUE DISTRIBUTIONS OF EDDINGTON RATIOS

In this section, we provide the physical explanation of the expected differences between “observed” and “true” distributions of M_{BH} and L/L_{Edd} . For a given physical parameter, we use “true” to denote its distribution of true values which is not actually observable because of measurement errors or uncertainties in adopted scaling relations. We use “observed” to denote the distribution of values obtained from the observations, by direct measurements or by applying scaling relations.

The classical version of the virial theorem which does not take into account radiation pressure provides a mass estimator which can be written as (Vestergaard & Peterson 2006)

$$M_{\text{BH}} = 10^f V_{\text{H}\beta}^2 L_{5100}^\gamma M_{\odot}, \quad (1)$$

where f is a scaling factor which encodes BLR geometry, physical structure, and projection effects, $V_{\text{H}\beta}$ is the FWHM of the broad H β line (units of 1000 km s^{-1}), and L_{5100} is the continuum luminosity at 5100 \AA (λL_{λ} , in units of $10^{44} \text{ erg s}^{-1}$). For consistency with N09 we adopt $f = 6.7$ and slope $\gamma = 0.6$.

When taking into account radiation forces due to the absorption of ionizing photons, the virial mass estimator is modified as follows (M08):

$$M_{\text{BH}} = 10^{f_{\text{rad}}} V_{\text{H}\beta}^2 L_{5100}^\gamma M_{\odot} + 10^{g'} L_{5100} M_{\odot}. \quad (2)$$

The second term represents the correction of radiation pressure and, in the assumption that BLR clouds are optically thick to ionizing photons and neglecting other sources of radiation pressure, it is given by (M08)

$$10^{g'} L_{5100} = \frac{L_{\text{ion}}}{4\pi G c m_p N_{\text{H}}} = \frac{b_{5100}/b_{\text{ion}}}{4\pi G c m_p N_{\text{H}}} L_{5100}, \quad (3)$$

where L_{ion} is the AGN luminosity in H-ionizing photons, N_{H} is the total column density of BLR clouds (of both ionized and neutral gas) toward the ionizing source, and b_{ion} and b_{5100} are the bolometric corrections for the ionizing and optical continuum luminosities, respectively ($L_{\text{bol}} = b_{\text{ion}} L_{\text{ion}} = b_{5100} L_{5100}$). The empirical calibration performed by M08, and adopted by N09, assumed $\gamma = 0.5$ and provides $f_{\text{rad}} \simeq 6.13$ and $g' \simeq 7.72$; the latter value indicates an average BLR column density $N_{\text{H}} = 10^{23} \text{ cm}^{-2}$, consistent with expectations from photoionization models. The column density of the BLR clouds sets the relative importance of the gravitational force (which depends on the cloud mass) and the radiative force (which is independent of mass). Therefore, it is the most critical parameter for the radiation pressure correction and we can outline its effects by writing

$$M_{\text{BH}} = 10^{f_{\text{rad}}} V_{\text{H}\beta}^2 L_{5100}^\gamma M_{\odot} + 10^g \frac{L_{5100}}{N_{23}} M_{\odot}, \quad (4)$$

where N_{23} is N_{H} in units of 10^{23} cm^{-2} . The Eddington ratios with and without radiation pressure correction are therefore

$$\left(\frac{L_{\text{bol}}}{L_{\text{Edd}}}\right)_{\text{vir}} = b_{5100} [L_{\text{Edd},\odot} 10^f V_{\text{H}\beta}^2 L_{5100}^{\gamma-1}]^{-1}, \quad (5)$$

$$\left(\frac{L_{\text{bol}}}{L_{\text{Edd}}}\right)_{\text{rad}} = b_{5100} \left[L_{\text{Edd},\odot} \left(10^{f_{\text{rad}}} V_{\text{H}\beta}^2 L_{5100}^{\gamma-1} + \frac{10^g}{N_{23}} \right) \right]^{-1}, \quad (6)$$

where $L_{\text{Edd},\odot}$ is the Eddington luminosity for a $1 M_{\odot}$ object. The observed distribution of Eddington ratios in the pure virial case is determined by the observed distributions of $V_{\text{H}\beta}$ and L_{5100} . For instance, if $V_{\text{H}\beta}$ and L_{5100} are log-normally distributed, then so are the Eddington ratios. Similarly, the observed distribution of Eddington ratios with the radiation pressure correction is determined by the observed distributions of $V_{\text{H}\beta}$ and L_{5100} but with an important additional feature which results from the addition of the radiation pressure correction. When the radiation pressure correction dominates over the virial term, $10^{f_{\text{rad}}-g} V_{\text{H}\beta}^2 L_{5100}^{\gamma-1} N_{23} \ll 1$, the Eddington ratio becomes asymptotically constant

$$\left(\frac{L_{\text{bol}}}{L_{\text{Edd}}}\right)_{\text{rad}} \longrightarrow \frac{L_{\text{BLR}}(N_{\text{H}})}{L_{\text{Edd}}} = b_{5100} \left[L_{\text{Edd},\odot} \frac{10^g}{N_{23}} \right]^{-1}, \quad (7)$$

where $L_{\text{BLR}}(N_{\text{H}})$ is the critical luminosity at which radiation forces on BLR clouds balance gravitation, and is a function of N_{H} ($L_{\text{BLR}} \propto N_{\text{H}}$). The physical meaning of this behavior is

straightforward: the virial mass estimator with the correction for the radiation pressure is valid only if the BLR is gravitationally bound, that is if $L_{\text{bol}} < L_{\text{BLR}}(N_{\text{H}}) < L_{\text{Edd}}$. Since $L_{\text{BLR}}(N_{\text{H}})$ is always smaller than L_{Edd} , $L_{\text{BLR}}(N_{\text{H}})/L_{\text{Edd}}$ is the maximum allowed Eddington ratio and any L/L_{Edd} distribution derived from Equation (4) will be truncated at that value.

At first sight, it might be assumed that the observed distributions of $L_{\text{bol}}/L_{\text{Edd}}$, obtained from Equations (5) and (6) and depending only on the observed distributions of $V_{H\beta}$ and L_{5100} , provide a good description of the “true” ones. However, this proposition does not take into account the fact that the scaling parameters themselves (f , f_{rad} , g , N_{23} , b_{5100}) are drawn from their own true distribution functions and necessarily vary from one object to another. In order to properly represent the true $L_{\text{bol}}/L_{\text{Edd}}$ distribution this dispersion must be included, and the interpretation of the observed differences between true and observed M_{BH} and $L_{\text{bol}}/L_{\text{Edd}}$ distributions must be modified accordingly. Since the scatter in the scaling parameters is not known and cannot be taken into account, the observed $L_{\text{bol}}/L_{\text{Edd}}$ distribution in the case of the radiation pressure correction will always present a sharp cutoff at $L_{\text{BLR}}/L_{\text{Edd}}$. This cutoff is smeared away when taking into account the true distribution of scaling parameters in general, and of BLR column densities in particular. Onken et al. (2004) showed that the rms scatter of the $M_{\text{BH}}-\sigma$ relation using virial masses is $\simeq 0.5$ dex, compared to the $\simeq 0.3$ dex of the same relation using more direct BH mass measurements from spatially resolved stellar or gas kinematics. Part of this additional scatter is probably explained by a broad distribution of f (or f_{rad}) values, which is naturally expected since the physical properties of BLR clouds must be characterized by a variance from one object to another (e.g., different cloud geometries and spatial distributions, relative orientations of the line of sights, etc.). Recently, fast eclipsing of the X-ray emitting source in the Seyfert galaxy NGC 1365 has been unambiguously explained by occultation from fast moving BLR clouds (Risaliti et al. 2007, 2009). In particular, these fast eclipses allow us to estimate the column density of BLR clouds toward the AGN, N_{H} finding a distribution of values in the $10^{23}-10^{24} \text{ cm}^{-2}$ range. The existence of a relatively broad distribution of N_{H} values in a single object might indicate an even broader distribution of values over the whole AGN population.

The L/L_{Edd} distribution plotted in Figure 3 of N09 clearly shows the features discussed above. In particular, the distribution obtained with the correction for radiation pressure is sharply cutoff at $L/L_{\text{Edd}} \simeq 0.15$ which corresponds to the critical value at which the BLR becomes gravitationally unbound. Conversely, the L/L_{Edd} distributions for type-2 and type-1 AGNs with classical virial masses are bell-shaped, with a tail beyond $L/L_{\text{Edd}} \simeq 0.15$. N09 interpreted these differences in the L/L_{Edd} distributions as an indication that BLR clouds have column densities significantly larger than 10^{23} cm^{-2} , for which the radiation-pressure force term is negligible. However, a distribution of relatively small N_{H} values ($\sim 10^{23} \text{ cm}^{-2}$) with a tail extending to large values could in principle provide an alternative explanation of the existence of type-2 AGNs with L/L_{Edd} beyond the critical value, without implying that radiation forces are not important in general.

3. MONTE CARLO SIMULATIONS

In this section, we use Monte Carlo simulations to analyze the discrepancies observed by N09 and test the importance of radiation forces. We start from the observed distributions of continuum luminosity and broad line widths which, combined

with assumed “real” distributions of bolometric correction and BLR column densities, allow us to obtain the real distribution of BH mass values and Eddington ratios. We consider two cases in which M_{BH} is obtained with or without important radiation pressure effect. By taking into account the intrinsic scatter in the $M_{\text{BH}}-\sigma$ relation, we then estimate the observed distribution of stellar velocity dispersion from M_{BH} . Then we repeat the analysis by N09 and compare the “observed” distributions of BH mass and Eddington ratios derived from the different scaling relations for M_{BH} .

For simplicity, we assume that the distributions of observed quantities and scaling parameters are log-normal. In particular, we assume that the observed distributions of L_{5100} and $V_{H\beta}$ values can be expressed as

$$\log(L_{5100}) = \mathcal{L} = \mathcal{L}_0 + \Sigma_{\mathcal{L}} i, \quad (8)$$

$$\log(V_{H\beta}) = \mathcal{V} = \mathcal{V}_0 + \Sigma_{\mathcal{V}} j, \quad (9)$$

where i and j are normally distributed random numbers with zero average and unitary standard deviation. \mathcal{L}_0 and \mathcal{V}_0 are therefore the averages of the distributions of $\log(L_{5100})$ and $\log(V_{H\beta})$ values, while $\Sigma_{\mathcal{L}}$ and $\Sigma_{\mathcal{V}}$ are the standard deviations. After the selection of the i and j random numbers, we apply the sample selection criteria adopted by N09 ($42.8 \leq \mathcal{L} + 44 \leq 44.8$ and $\mathcal{V} > 1$). The AGNs are characterized by true distributions of bolometric corrections and BLR cloud column densities:

$$\log(b_{5100}) = \mathcal{B} = \mathcal{B}_0 + \Sigma_{\mathcal{B}} h, \quad (10)$$

$$\log(N_{\text{H}}) = \mathcal{N} = \mathcal{N}_0 + \Sigma_{\mathcal{N}} k, \quad (11)$$

with the same conventions as above. It is then possible to recover the true distribution of BH masses from the \mathcal{L} , \mathcal{V} , \mathcal{B} , and \mathcal{N} distributions. When radiation forces are important,

$$\log(M_{\text{BH}}/M_{\odot}) = \mathcal{M}_{\text{R}} = \log(10^{\mathcal{M}_1} + 10^{\mathcal{M}_2}), \quad (12)$$

with

$$\mathcal{M}_1 = (f_{\text{rad}} + \Sigma_f u) + 2\mathcal{V} + \gamma\mathcal{L} \quad (13)$$

$$\mathcal{M}_2 = g + \mathcal{L} + (23 - \mathcal{N}), \quad (14)$$

where $\Sigma_f u$ represents the variance of the f_{rad} scaling factor and we have assumed, for simplicity, that the variance in g is entirely dominated by the variance in the column density. Conversely, in the case where radiation forces are not important, the true distribution of BH masses is given by

$$\mathcal{M}_{\text{V}} = (f + \Sigma_f u) + 2\mathcal{V} + \gamma\mathcal{L}. \quad (15)$$

The observed distribution of stellar velocity dispersions is obtained, in both cases, from the $M_{\text{BH}}-\sigma$ relation:

$$\mathcal{M}_{\text{R}} \text{ (or } \mathcal{M}_{\text{V}}) = \alpha + \beta\mathcal{S} + \Sigma_{\mathcal{M}} w, \quad (16)$$

where $\mathcal{S} = \log(\sigma/200 \text{ km s}^{-1})$ and w is a normally distributed random number like i . Thus, it is possible to derive the “observed” BH masses and bolometric luminosities as

$$\bar{\mathcal{M}}_{\sigma} = \mathcal{M}_{\text{rad}} \text{ (or } \mathcal{M}_{\text{vir}}) - \Sigma_{\mathcal{M}} w, \quad (17)$$

$$\bar{\mathcal{M}}_{\text{vir}} = f + 2\mathcal{V} + \gamma\mathcal{L}, \quad (18)$$

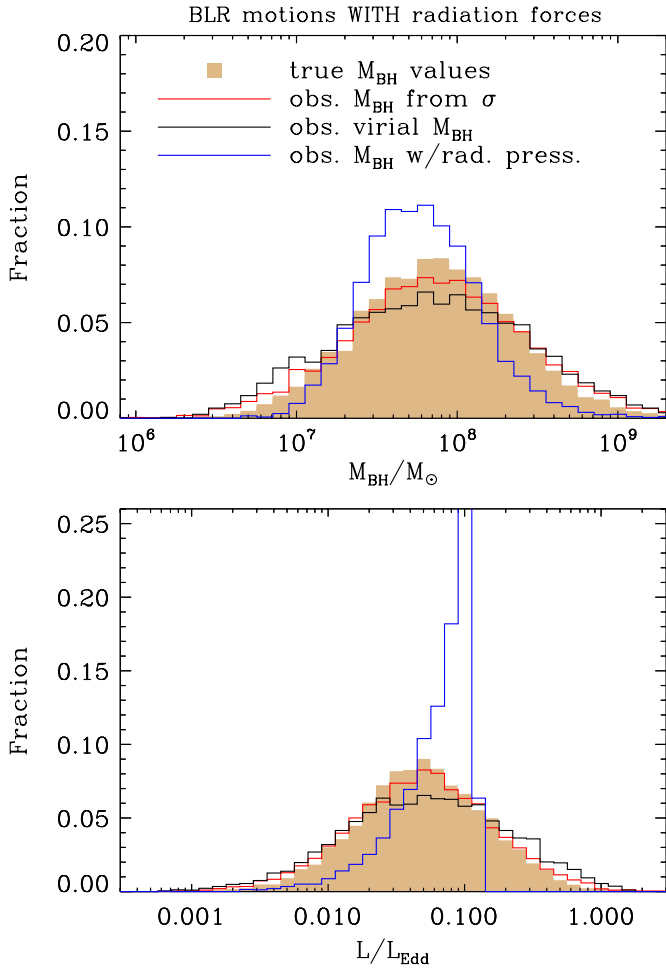


Figure 1. Observed and true distributions of BH masses (top panel) and Eddington ratios (bottom panel) in the case of BLR motions affected by radiation forces. The red line denotes BH masses derived from the $M_{\text{BH}}-\sigma$ relation (\mathcal{M}_{σ}), the black line denotes BH masses computed using classical virial relations (\mathcal{M}_{vir}), and the blue line denotes BH masses computed using virial relations corrected for radiation pressure (\mathcal{M}_{rad}). The shaded histograms indicate the true distributions of BH masses and Eddington ratios. The sharp cutoff in the observed L/L_{Edd} distributions is present even if radiation pressure forces are important.

$$\bar{\mathcal{M}}_{\text{rad}} = \log(10^{f_{\text{rad}}+2\nu+\gamma\mathcal{L}} + 10^{g+\mathcal{L}}), \quad (19)$$

$$\bar{\mathcal{L}}_{\text{bol}} = \mathcal{B}_0 + \mathcal{L}, \quad (20)$$

where $\bar{\mathcal{M}}_{\text{rad}}$ or $\bar{\mathcal{M}}_{\text{vir}}$ represent the virial estimators with or without correction for radiation pressure. These provide the observed distributions of L/L_{Edd} and, clearly, these observational quantities do not include the true distributions of scaling parameter values.

In Figure 1, we show the results for the case of BLR motions affected by radiation forces. We use the following set of values: 5000 random realizations with $(\mathcal{L}_0, \Sigma_{\mathcal{L}}) = (-0.2, 0.3)$, $(\mathcal{V}_0, \Sigma_{\mathcal{V}}) = (0.65, 0.3)$, $(\mathcal{B}_0, \Sigma_{\mathcal{B}}) = (0.95, 0.2)$, $(\mathcal{N}_0, \Sigma_{\mathcal{N}}) = (23.0, 0.5)$, $\Sigma_f = 0.3$. This set of parameters was chosen to reproduce the \mathcal{M}_{σ} and $\bar{\mathcal{M}}_{\text{vir}}$ distributions observed by N09 but the actual adopted values do not influence our general conclusions. Our simulated samples nicely reproduce the features of the $\bar{\mathcal{M}}_{\text{rad}}$ and L/L_{Edd} distributions observed by N09: the distribution of BH masses computed with the radiation pressure correction is narrower than that based on $M_{\text{BH}}-\sigma$ and classical virial relations; the observed distributions of L/L_{Edd} ratios based on $M_{\text{BH}}-\sigma$ and classical virial relations are in nice

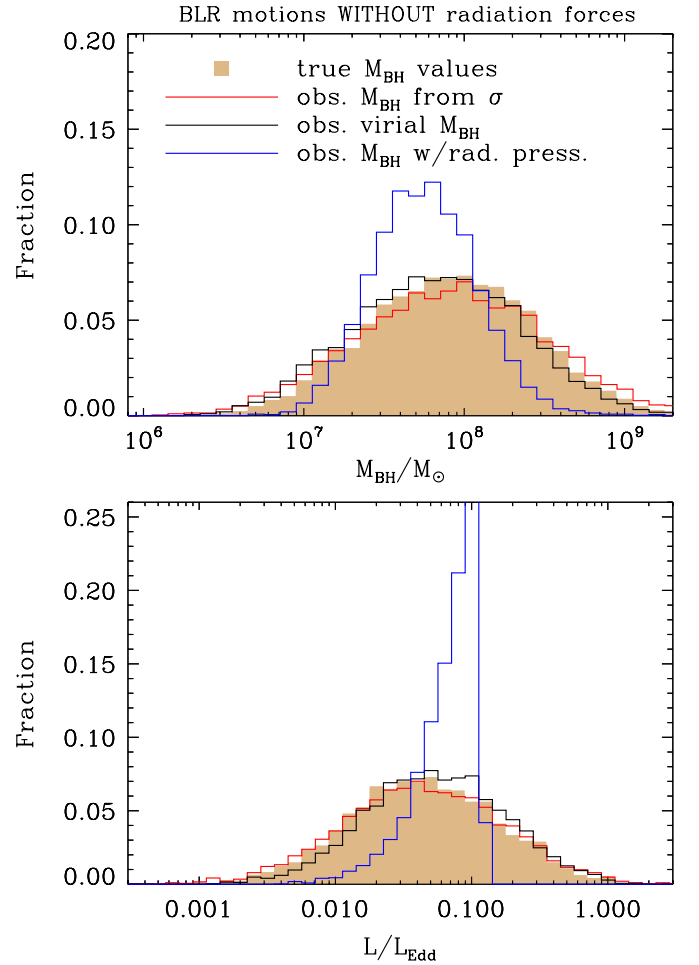


Figure 2. Observed and true distributions of BH masses (top panel) and Eddington ratios (bottom panel) in the case of BLR motions NOT affected by radiation forces. Notation as in Figure 1.

agreement, while that based on radiation-pressure-corrected virial masses is sharply truncated at $L/L_{\text{Edd}} \simeq 0.15$ and more sharply peaked. The true M_{BH} and L/L_{Edd} distributions are intermediate between those found with and without the radiation pressure correction. For this simulation, we have adopted $\Sigma_{\mathcal{N}} = 0.5$ dex but if we increase that value the true M_{BH} and L/L_{Edd} distributions will be broader and approaching those observed when using M_{BH} without the radiation pressure correction. Allowing for an intrinsic dispersion in the scaling parameters, especially in N_{H} , smears the cutoff in the observed L/L_{Edd} distribution which then approaches the true one. The comparison between the observed (blue line) and true (shaded histogram) L/L_{Edd} distributions shows that a fraction of sources have Eddington ratios larger than the critical value at which a BLR with $N_{\text{H}} = 10^{23} \text{ cm}^{-2}$ becomes gravitationally unbound. This is possible only because, allowing for a distribution of N_{H} values, a fraction of the sources have $N_{\text{H}} > 10^{23} \text{ cm}^{-2}$. Then, if the true distribution of Eddington ratios extends to large values, BLRs are gravitationally bound only if their average column density increases with L/L_{Edd} (see also Dong et al. 2009).

Finally, in Figure 2 we show the results for the case of BLR motions NOT affected by radiation forces. We use the same set of parameters as before, except for $\Sigma_{\mathcal{V}} = 0.25$ to have similar true distributions of M_{BH} and L/L_{Edd} . We observe again the same features that prompted N09 to consider radiation forces unimportant, and in particular the sharp cutoff in the observed

L/L_{Edd} distribution. In this case, the true distributions of M_{BH} and L/L_{Edd} values are well matched by the observed ones, because the intrinsic scatter of the scaling parameters is small compared to the combined scatter of the observed line widths and continuum luminosities.

4. SUMMARY

The simulations presented in Figures 1 and 2 indicate that, when taken at face value, the comparison of the observed M_{BH} and L/L_{Edd} distributions would always lead to the conclusions that radiation forces are not important in determining the motions of BLR clouds. However, we have shown that it is not possible to assess the importance of radiation forces on BLR cloud motions on the basis of the *observed* M_{BH} and L/L_{Edd} distributions. The differences between the L/L_{Edd} distribution obtained using the $M_{\text{BH}}-\sigma$ relation and the one based on virial masses with radiation pressure correction can also be explained by neglecting the intrinsic dispersion in the adopted scaling parameters (e.g., f_{rad} and g). In particular, the sharp cutoff at $L/L_{\text{Edd}} \simeq 0.15$ observed by N09 corresponds to the critical luminosity at which radiation forces balance gravitational attraction on BLR clouds with the adopted N_{H} value (10^{23} cm^{-3}) but does not indicate that radiation forces are negligible. A broad distribution of N_{H} values (e.g., log-normal with 10^{23} cm^{-2} average and 0.5 dex standard deviation) will remove such sharp cutoff, fully explaining the observed differences in L/L_{Edd} distributions. The sources with $L/L_{\text{Edd}} > 0.15$ are then those in the high tail of the N_{H} distribution. Moreover, at these large L/L_{Edd} , BLR clouds can be gravitationally bound only if they have large column densities, e.g., if N_{H} increases with L/L_{Edd} . In this scenario, the sources with $N_{\text{H}} \simeq 10^{24} \text{ cm}^{-3}$, for which radiation pressure is negligible, constitute a minority of the whole population. In conclusion, it is not possible to distinguish between the two scenarios in which radiation forces are negligible in all sources or in just a minority of them.

When estimating BH masses using virial mass estimators, one should then always consider the possibility of important radiation forces by using the currently calibrated correction which corresponds to $N_{\text{H}} \simeq 10^{23} \text{ cm}^{-2}$ (M08), until it is possible to assess the possible dependence of N_{H} on the observed source properties.

Finally, the comparison of the M_{BH} and L/L_{Edd} distributions proposed by N09 can still constrain the overall scatter in the scaling parameters used in virial mass estimators. A detailed analysis of this issue is beyond the scope of this Letter but, briefly, after obtaining the true distributions of M_{BH} and L/L_{Edd} from the observed distribution of stellar velocity dispersions one can match it with the M_{BH} and L/L_{Edd} distributions based on virial masses, convolved with a suitable broadening function. Such a broadening function would then provide the combined variance of BLR physical properties in the sample of AGNs under examination.

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